## FIVE INNOVATIONS IN PLANETARY THEORY

## INTRODUCED BY KEPLER IN ASTRONOMIA NOVA

1. Lines of apsides -- i.e. lines between aphelia and perihelia -- always pass through the true Sun, and all motions should be referenced to the true Sun rather than the mean Sun.
2. Introduce a bi-section of the eccentricity for the Earth-Sun orbit, in the manner of the other orbits, implying that the Earth-Sun motion involves a true variation in velocity, not just an apparent one; the key consequence of bi-section of the eccentricity is that, with the equant or any velocity rule refining it, the arc length velocities of the orbiting body at its maximum and minimum distances from the central body -- i.e. at aphelion and perihelion when the Sun is the central body -- are in the inverse ratio of these distances.
3. Replace the equant with the area rule -- or its equivalent, the component of the arc length velocity normal to a radius drawn from the central body varies inversely with the distance from the central body -- for purposes of determining the (unperturbed) location of the orbiting body versus time from aphelion.
4. Construe the trajectory of the orbiting body (or, at least, its primary motion at any time) as defining an ellipse with the central body at a focus.
5. Construe the orbital motion (at least at any time) as occurring in a plane that passes through the true Sun and is inclined at a measurable angle to the plane defined by the Earth-Sun orbit.

Qualifying remarks in parentheses in the above statements represent allowances Kepler expressly made in subsequent works; they are not inconsistent with the conclusions of Astronomia Nova, for what they amount to, in effect, is to allow for the possibility of the values of the elements of the Keplerian orbits changing over time. All five of these innovations are still part of orbital astronomy today.


Fig. 5. Errors in geocentric longitudes of Mars. (Above) Ptolemaic calculation compared with Kepler's Rudolphine tables; (below, at much enlarged scale) the Rudolphine errors compared with those from Magini's "Keplerian" method in his Supplementum ephemeridum.

| Tempus | Locus $\odot$ | $\begin{gathered} \text { Solis a } \\ \text { Terra } \\ \text { distantia } \\ \hline \end{gathered}$ |  | Martis eccentricus in ecliptica | Locus computatus | Locus observatus | Differentia | Lati－ tudo |
| :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: | :---: |
| $\begin{aligned} \text { 1582. } 23 & \text { Nove. H. } 16 . \quad 0 \\ 26 & \text { Dece. } \end{aligned} \text { H. 8. 8. } 30$ | $\begin{aligned} & 11^{\circ} \cdot 41^{\prime} x^{x} \\ & 15 . \\ & 196 \\ & 19 . \\ & 16 . \\ & 163 \\ & \hline 6 \end{aligned}$ | $\begin{aligned} & 98345 \\ & 98226 \\ & 98222 \\ & 98624 \end{aligned}$ | $\begin{aligned} & 158852 \\ & 162104 \\ & 162443 \\ & 164421 \end{aligned}$ |  | 26． 40.059 <br> 17．44． 196 <br> 8．17． 576 | 26．38． 30 6 <br> 17． 40.3069 <br> 16． 0.3069 <br> 8．20． 3069 | $\begin{aligned} & 1^{\prime} \cdot 30^{\prime \prime}+ \\ & 3 \cdot 49 \\ & 5 \cdot \\ & 50 \\ & 2 \cdot \\ & 2 . \end{aligned} 33-1$ | Bor． <br> 2.49 <br> 4． 7 <br> 2． 52 |
| 1584． 21 Dece．H． $14 . \quad \circ$ 1585． 24 Janua．H． 4 Febr． H． 12 F Mart． H． 40 | 10． 167 <br> 14． 53 m <br> 26． 10 ※ <br> 2． 16 V | $\begin{aligned} & 98207 \\ & 98595 \\ & 98830 \\ & 99898 \\ & \hline \end{aligned}$ | 164907 166400 <br> 166170 | $\begin{array}{ccccc} 3 . & 51 . & 45 \\ 18 . & 47 & 8 & \Omega \\ \text { 23. } & 33 & 83 & 41 & \Omega \\ 9 . & 23 . & 14 & \mathrm{ml} \end{array}$ | $\begin{aligned} & 1.14 .34 \mathrm{mp} \\ & \text { 24. } 3.58 \Omega \\ & \text { 19. } 43.52 \Omega \\ & 11.43 \cdot 31 \Omega \\ & \hline \end{aligned}$ | 1． 13.30 Il <br> 24． $7.30 \Omega$ <br> 19．47．$\circ$ o $\delta$ <br> 11．46．$\circ \Omega$ | 1． $4+$ <br> 3． 32 <br> 3． 8 <br> 2． 29 | $\begin{aligned} & 3 \cdot 31 \\ & 4.31 \\ & 4.28 \end{aligned}$ $3.22$ |
| 1587. 25 Janua． H． 17. <br> 4 Mart． H． 13.24  <br> 10 Mart． H． 11.30  <br> 21 April． H． 9. | 16． 1 m <br> 24．$\circ x$ <br> 29． 52 K <br> 10． $4^{8}$ y | $\begin{aligned} & 98611 \\ & 99595 \\ & 99780 \end{aligned}$ $101010$ | $\begin{aligned} & 166232 \\ & 164737 \\ & 164382 \\ & 161027 \end{aligned}$ | 8．13． 40 m <br> 24．56．50 M11 <br> 27．35． 54 III <br> 16． $44.51 \bumpeq$ | 4．41． $50 \bumpeq$ 26．24． $4^{1} \mathrm{ny}$ 24． 5.15 mp 15．49． 50 mp | $\begin{aligned} & \text { 4. 42. } 0 \bumpeq \\ & \text { 26. 25. } 40 \mathrm{mp} \\ & \text { 24. 5. } 5 \mathrm{mp} \\ & \text { 15. } 48.20 \mathrm{mp} \end{aligned}$ | $\begin{array}{lll} \text { o. } & 10 & - \\ \text { o. } & 59 & - \\ \text { o. } & 0 & - \\ \text { 1. } & 30 & + \end{array}$ | $\begin{aligned} & 3.26 \\ & 3.38 \\ & 3.29 \\ & 1.48 \end{aligned}$ |
| 1589． 8 Mart．H． 16.24  <br> 13 April． <br> H． 11.15  <br> 15 April． H．12． 5 <br> 6 Maji． <br> H．11． 20  | $\begin{aligned} \text { 28. } & 36 x \\ 3 . & 38 y \\ 5 . & 36 y \\ 25 . & 498 \end{aligned}$ | 99736 100810 100866 | $\begin{aligned} & 161000 \\ & 157141 \\ & 156900 \\ & 154326 \end{aligned}$ | $\begin{array}{ccccc} \text { 16. } & 55 & 1 & 1 & \imath \\ 4 . & 1 . & 50 & \mathrm{~m} \\ 5 . & 1 . & 41 & \mathrm{ml} \\ 15 . & 30 . & 36 & \mathrm{ml} \end{array}$ | $\begin{array}{r} \text { 12. 14. } 7 \mathrm{ml} \\ \text { 4. } 45 . \circ \mathrm{m} \\ \text { 3. } 58.57 \mathrm{ml} \\ \text { 27. } 8.17 \bumpeq \end{array}$ | 12．16． 50 m <br> 4． 43.20 m <br> 3． 58.20 m <br> 27． $7.20 \bumpeq$ | $\begin{aligned} & \text { 2. } 43 \\ & \text { 1. } 40+ \\ & \text { o. } 37+ \\ & \text { o. } 57+ \end{aligned}$ | $\begin{aligned} & \text { 2. } 4 \\ & \text { 1. } 10 \\ & \text { 1. } 4 \\ & \text { o. } 7 \\ & \hline \end{aligned}$ |
| $\begin{array}{rlll} \text { 1591. } 13 \text { Maji. } & \text { H. 14. } 0 \\ \text { 6 Junii } & \text { H. 12. } 20 \\ \text { 10 Junii } & \text { H. 11. } 50 \\ 28 \text { Junii } & \text { H. 10. } 24 \\ \hline \end{array}$ | $\begin{aligned} & \text { 2. } 10 \text { 11 } \\ & \text { 24. } 59 \mathrm{I} \\ & \text { 18. } 47 \mathrm{I} \\ & \text { 15. } 51 \end{aligned}$ | $\begin{aligned} & 101467 \\ & 101769 \\ & 101789 \\ & 101770 \end{aligned}$ | $\left\|\begin{array}{l} 147891 \\ 144981 \\ 144526 \\ 142608 \end{array}\right\|$ | $\begin{array}{lllll} 12 . & 7 \cdot & 38 & 7 \\ \text { 25. } & 38 \cdot & 48 & x^{7} \\ \text { 27. } & 56 \cdot & 49 & \boxed{x} \\ \text { 8. } & 29 \cdot & 32 & 7 \end{array}$ | 2． 19.36 万 <br> 27．11．45 㐅 <br> 25．57． 57 ス $^{\text {x }}$ <br> 21．4． $21 \not \chi^{\star}$ |  | $\begin{aligned} & \text { 4. } 24 \\ & \text { 3. } \\ & \text { 4. } \\ & \text { 5. } \\ & \text { 5. } \end{aligned}$ | Aust． <br> 2． 25 <br> 3． 55 <br> 4． 8 <br> 4． 45 |
| 1593.21 Julii H．14．o  <br> 22 Aug． H．12． 20  <br> 29 Aug． H．10． 20  <br> 3 Octo． H．8． o | $\begin{array}{r} \text { 8. } 26 \Omega \\ \text { 9. } 11 \mathrm{~m} \% \\ \text { 11. } 54 \mathrm{mp} \\ \text { 20. } 15 \bumpeq \end{array}$ | $\begin{array}{r} 101498 \\ 100761 \\ 100562 \\ 99500 \\ \hline \end{array}$ | $\begin{aligned} & 138376 \\ & 138463 \\ & 138682 \\ & 140697 \end{aligned}$ |  | $\begin{aligned} & \text { 17. 43. } 14 \mathrm{~K} \\ & \text { 13. } 9.39 \mathrm{~K} \\ & \text { 11. 11. } 41 \mathrm{~K} \\ & \text { 7. } 49.54 \mathrm{~K} \\ & \hline \end{aligned}$ | 17． 45.45 K <br> 13．10．is $X$ <br> 11．14．$\circ x$ <br> 7． 50.10 x | 2. <br> 31 $\square$ <br> o． 3 $\qquad$ <br> 2． 19 <br> o． 16 $\qquad$ $\qquad$ | 5.46 6.7 5.52 3.17 |
| 1595. 17 Sept． H．16． 45 <br> 27 Octo． H．12． 20  <br>     <br> 3 Nove．H．12． $\circ$  <br> 18 Dece． H．8． $\circ$ | 4． $18 \bumpeq$ <br> 13． 59 m <br> 21． 2 m <br> 6． $43 \pi$ | 98851 <br> 98694 <br> 98200 | $\begin{aligned} & 143222 \\ & 147^{890} \\ & 148773 \\ & 154539 \end{aligned}$ | $\begin{array}{cccc} \text { 22. } & 49 . & 19 & 8 \\ \text { 15. } & 35 & 38 & 8 \\ & & & 8 \\ \text { 19. } & 26 . & 33 & 8 \\ \text { 13. } & 2 . & 29 & \boxed{I I} \end{array}$ | $\begin{array}{llll} \text { 26. } & 5 . & 45 & 8 \\ \text { 18. } & 50 & 46 & 8 \\ \text { 16. } & \text { 18. } 33 & 8 \\ \text { 11. } & 39 . & 1 & 8 \end{array}$ | 26．7． 128 <br> 18．51．15 8 <br> 16．18． 30 y <br> 11．40． 0 8 | 1． 27 － <br> o． 29 － <br> o． $3+$ <br> o． $59-$ | $\begin{aligned} & 1.42 \\ & \text { o. } 6 \\ & \text { Bor. } \\ & \text { o. } 17 \\ & 1.40 \end{aligned}$ |

## ASTRONOMIA NOVA OVERVIEW

## GOALS

1. Introduce reforms to orbital theory for Mars to remove all excessive discrepancies vs. observation
2. Determine true motion of Mars at least with respect to the Sun, if not the fixed stars, for 1580-1604
3. Establish Copernican theory for Mars-Earth-Sun

## PROBLEMS

1. Working from a limited set of observations, 15801604, restricted to geocentric longitudes \& latitudes
2. Must use theory to reach any conclusions re orbital motion, but without begging questions re the motion
3. Observations were of limited precision (though with estimated bounds of the limits), and corrections for parallax and atmospheric refraction were suspect
4. Supplemental theories, e.g. of Earth-Sun orbit, can introduce extraneous sources of discrepancy

## SUPPLEMENTAL THEORIES EMPLOYED

1. Tycho's Earth-Sun theory, along the way refined to incorporate bisection of eccentricity
2. Vicarious hypothesis for heliocentric longitudes of Mars about the true Sun
3. (In places) Tycho's theory of heliocentric longitudes for Mars about the mean Sun
4. Assumption that Mars orbits the Sun, not the Earth (while not foreclosing a Ptolemaic representation)

## Questions of Evidence

To what extent did the evidence presented in Astronomia Nova on the Mars and Earth-Sun orbits exclude alternatives to:

1. The line of apsides passes through the true Sun?
2. The orbit lies on a plane passing through the true Sun at a constant angle of inclination with respect to the ecliptic?
3. The bisection of the eccentricity of the EarthSun orbit?
4. The diametral distance rule?
5. The area rule?
6. Mars's trajectory is an ellipse?

To what extent did the evidence presented in Astronomia Nova support the Copernican system:

1. Over the Tychonic system?
2. Over the Ptolemaic system?

## Why Kepler?

$>$ Access to Tycho's observations
$>$ Tycho's redefinition of the goal of astronomy: agreement within observational accuracy at all times around the zodiac
$>$ Kepler's demonstration of limitations in relying on acronychal observations
$>$ Idea, inherited from Copernicus, of employing theory-mediated triangulations to determine sun-planet and earth-sun distances vs. time
$>$ Tycho's exceptionally accurate theory of the Sun
$>$ The (perhaps coincidental) choice of Mars


## December 27, 1571

Born in Weil der Stadt, Germany

## 1584-1588

Attends seminary school at Adelberg and Maulbronn

## 1589-1594

Attends the University of Tübingen, where he receives a
B.A. (by examination, 1588) and M.A. (1591); nearly com-
pletes three additional years of study in theology

## April 1594

Arrives in Graz, Styria, to assume the position of mathematics teacher and district mathematician

## 1596-March 1597

Printing of the Mysterium cosmographicum in Tübingen

## April 27, 1597

Marries Barbara Müller

## September 28, 1598

Counter-Reformation begins in Styria; Protestant teachers and preachers expelled from Graz; Kepler is allowed to return after about a month

## January-June 1600

Visits Tycho Brahe at Benatky Castle

## September 30, 1600

Leaves Styria with his family and all their possessions
when all remaining Protestants are banished

## October 24, 1601

Tycho Brahe dies; two days later, Kepler is named his successor as imperial mathematician to Rudolf II in Prague

## 1604

Publishes the Astronomia pars optica

## Around Easter, 1605

Discovers the elliptical form of Mars's orbit

## 1609

Astronomia nova finally published

## March 1610

Galileo publishes the Sidereus nuncius containing his telescopic discoveries; Kepler responds, publishing the Dissertatio cum nunceo sidereo in Prague in May

## Summer 1611

Kepler publishes Dioptrice, containing explanation of the telescope

July 3, 16 II
Barbara Kepler dies
January 20, 1612
Holy Roman Emperor Rudolf II dies; Archduke Matthias succeeds him

## May 1612

Begins work as mathematician to the Estates of Upper Austria in Linz

October 30, 1613
Marries Susanna Reuttinger
July 1615
Publishes Nova stereometria doliorum vinariorum in Linz

## Fall 1617

Publishes the first volume of Epitome astronomiae Copernicanae in Linz

## Fall 1617-early 1618

Returns to Württemberg with his mother, but her court case is delayed
May 15, 1618
Discovers his third law of planetary motion

## 1619

Harmonice mundi libri $V$ published in Linz
March 20, 1619
Holy Roman Emperor Matthias dies; Archduke Ferdinand II suceeds him 5 months later

## Early 1620

Second volume of Epitome astronomiae Copernicanae published in Linz

## August 7, 1620-August 1621

Katharina Kepler arrested for witchcraft; Johannes Kepler returns to Württemberg to assist in her defense

## Fall 1621

Final volume of the Epitome astronomiae Copernicanae published in Frankfurt

## October 1625

Counter-Reformation begins in Upper Austria

## November 1626

Kepler and his family leave Linz
December 1626-September 1627
Rudolfine Tables printed in Ulm

## July 1628

Kepler arrives in Sagan to become personal mathematician to General Wallenstein; Counter-Reformation in Sagan begins four months later

## November 15, 1630

Dies while visiting a meeting of imperial electors in Regensburg

May 23, 1618
"Defenestration of Prague"; Thirty Years War begins

This detail of the frontispiece of the Rudolfine Tables shows Kepler as the architect of the achievement. He works by candlelight on a model of the temple above him. The banner lists his important publicotions. A few coins dropped by the Hapsburg eagle have reached his desk


# Kepler's Principal Publications 

## Mysterium Cosmographicum (1596)

On More Certain Foundations of Astrology (1602)
Astronomiae Pars Optica Traditur (1604)
De Stella Nova (1606) - on the nova of 1604
Astronomia Nova (1609)
Dissertatio cum Nuncio siderio (1610)
Narratio de Jovis satellitibus (1611)
Dioptrice (1611) - theory of the telescope
Stereometria dolorioum vinariorum (1615)
Harmonice Mundi (1618)
Epitome Astronomiae Copernicanae (1618-1621)
De cometis libelli tres (1619) - on the comet of 1618
Mysterium Cosmographicum (revised, 1621)
Tabulae Rudolphinae (1627)
Admonitio ad Astronomos rerumque celesti studiosos, de mirisque rarisque anni 1631phaenomensis (1630)

Somnium seu astronomia lunari (1634)

After 1615, Ephemerides announcing astronomical phenomena on a regular basis, including one of 1630 above announcing 1631 transits of Mercury \& Venus

5.2. Kepler's nested polygons and planetary spheres (redrawn for greater clarity in the central portion), from the Mysterium cosmographicum.

Table 5.1 Relative sizes of the adjacent planetary orbits (assuming the innermost part of the outer orbit to be 1000)

| Planet | Intervening <br> regular solid | Computed <br> by Kepler | From <br> Copernicus |
| :--- | :--- | :--- | :--- |
| Saturn | Cube | 577 | 635 |
| Jupiter | Tetrahedron | 333 | 333 |
| Mars | Dodecahedron | 795 | 757 |
| Earth | Icosahedron | 795 | 794 |
| Venus | Octahedron | 707 | 723 |
| Mercury |  |  |  |

TABEILA ITT
ORBIVM PLANETARYM DIMENSIONES, ET DISTANTIAS PER QVINQVE REGVLARIA CORPORA GEOMETRICA EXHIBENS.


# Ioannis Keppleri HARMONICES M V N D I LIBRIV. Qvorvm 

Prinus Geometricvs, De Figurarum Regularium, que Proportio: nes Harmonicas conftituunt, ortu \& demonftrationibus.
Seciundus Architectonicvs; feu ex Gecmetria Figvrata, De Figuratum Regularium Cohgruentia in plano vel folido:
Tertius proprićHarmonicus, De Proportiontım Harmonicarum ortu ex Figuris; deque Natutâ \& Differentís rerum ad cantum pertinentium, contra Veteres:
Quarlus Metaphysicvs, Psychologicvs \& Àstrol ogicvs, De Harmoniarum mentali Eflentiâ earumque generibus in Mundo; prefertim de Harmonia radiorum, ex corporibus cocleftibus in Terram de(cendentibus, eiufque effectu in Natura feu Anima fublunary ic Humana:
Quintus Astronomičs \& \& Meł Aphyisicvs, De Harmoñiis abfolntifimis motuum cocleftium, ortuque Eccentricitatum ex proportionibus Harmonicis.
Appendix haber comparationem huius Operis cumi Harmonices Cl. Ptolemxi librollI I.cumque Roberci de Fluctibus,dictiFlud.Medici Oxonienfis fpeculationibus Harmonicis, operi de Mạcrocofmo. \& Microcofmo infertis.


Cum S.C. ovit Priuilegio adannos XV.
Lincii Auftrix,
Sumptibus Godofredi Tampachir Bibl. Francof. Excudebat Io annes Planc vs.

## KEPLER'S THIRD "LAW"

## FROM KEPLER'S DATA

|  | Period <br> (years) | $\begin{gathered} \text { Mean } \\ \text { Distance } \end{gathered}$ | Period <br> Squared | Distance Cubed |
| :---: | :---: | :---: | :---: | :---: |
| Mercury | 0.2408 | 0.3881 | 0.0580 | 0.0585 |
| Venus | 0.6152 | 0.7241 | 0.3785 | 0.3797 |
| Earth | 1.0000 | 1.0000 | 1.0000 | 1.0000 |
| Mars | 1.8808 | 1.5235 | 3.5376 | 3.5361 |
| Jupiter | 11.8621 | 5.200 | 140.7086 | 140.61 |
| Saturn | 29.4571 | 9.510 | 867.7198 | 860.08 |
|  | Period <br> (days) | a from Periods | a from Orbits | Percent Difference |
| Mercury | 87.97 | 0.3871 | 0.3881 | 0.26 |
| Venus | 224.70 | 0.7233 | 0.7241 | 0.11 |
| Earth | 365.25 | 1.0000 | 1.0000 | ---- |
| Mars | 686.98 | 1.5237 | 1.5235 | 0.01 |
| Jupiter | 4332.62 | 5.2012 | 5.200 | 0.02 |
| Saturn | 10759.20 | 9.5381 | 9.510 | 0.29 |

## Books IV and V of Kepler's Epitome: An Overview

"... all the more strange to them [readers] would be this Fourth Book, which airs so many new and unthought-of things concerning the whole nature of the heavens - so that you might doubt whether you were doing a part of physics or astronomy, unless you recognized that speculative astronomy is one whole part of physics." (p. 5)

## Book IV

Part I: System of the world, pp. 13-47 (of Prometheus ed.)

1. On the Principal Parts: p. 16, no solid spheres
2. On the Place of the Sun at the center: true Sun, p. 20
3. On the order of the movable spheres

With respect to the perfect solids p. 23ff
The harmonies: p. 31-32
4. On the ratios of the Principal Bodies to one Another on imperceptibility of annual stellar parallax: p. 43ff

Part II: On the Movement of Bodies: pp. 47-88

1. How many and what sort of movements denial of motion of fixed stars
2. Concerning the Causes of the Movements: pp. 48-55
"inertia" and the need for a motive force
3. On the Revolution of the Sun about its Axis, and Its Effects: pp. 55-65
solar magnetic emanation: pushing planets
strength varies inversely with distance
p. 63: why not as inverse-square, like light
4. On the Causes of the $3 / 2$ Power Ratio: pp. 65-67
5. On the Annual Movement of the Earth: pp. 67-88
p. 67: five questions
p. 68: appeal to phases of Venus
p. 70: eight reasons for true Sun
pp. 71-76: eighteen reasons to reject Tychonic
pp. 78-88: rotation of Earth, Jupiter; Moons
Part III: On the Real and True Irregularity of the Planets and its Causes (pp. 88120)
6. The causes of the true irregularities: pp. 88-93
three part answer: p. 89
7. On the Causes of Irregularity in Longitude: pp. 93-94
laws of variation in speed: lever, balance analogy
8. Causes of Irregularity in Altitude (i.e. $r$ or $S P$ ): pp. 95-106
p. 98: no interaction between planet and Sun
p. 99: basic account
pp. 101-105: thread displacement, restoration; precession of apsides
9. On the Movement of the Latitude: pp. 106-112
basic answer: pp. 106-107
pp. 107-112: why inclination at all; why different from one to another
10. On the Twofold Irregularities of the Moon and Their Causes: pp. 112-120 basic answer: pp. 113-114
precession of the line of apsides: p. 117

Book V (pp. 124-164)
Part I: On the Eccentric Circles or Schemata of the Planets (pp. 124-146)
p. 125: basic answer (figure)
p. 127: seven tasks

1. Concerning the Increment in Libration: p. 128
why sinusoidal (p. 129); balance (p. 131)
2. On the Sum of the Libration Gone Through With (p. 133)
(i.e. the cumulative or progressive change in $r$ )
3. On the Figure of the Orbit: pp. 135-139
i.e. circumscribed circle, eccentric anomaly, and varying $r$
4. On the Measure of Time (with apology for statement in Astronomia Nova): pp. 139-143
5. On the Equivalence of the Circle of the Plane and the Plane of the Ellipse in Measuring Delays: pp. 143-145
"Kepler's problem"
6. On the Regularity of the Digression in Latitude: pp. 145-146
i.e. motion of the nodes

Part II: On the Astronomical Terms Arising From Calculations and the Eccentric Orbit (pp. 146-164)
(an explanation of important terms)

1. Concerning Designation
2. Concerning Libration
3. Concerning the Delay of the Planet in any Arc
4. Concerning the Angle of the Sun
5. On the Digression of the Planets Away from the Ecliptic

## 6. On the Movement of the Apsides and Nodes

For the contents of the first three books, see p. 4; for the contents of Books VI, on the individual orbits, and VII, see p. 164; the following is from Book VI:

|  |  |  |  | Eccentricitates qualium semidiameter est 100000 . |
| :---: | :---: | :---: | :---: | :---: |
|  | Aphelium. | Medium. | Perihelium. |  |
| Saturni . | 100 207. | 951000. | 896793. | 5700 |
| Iouis | 544708. | 519650. | 494592. | 4822 |
| Martis | 166465. | 152350. | 138235. | 9263 |
| Orbis Magni | 101800. | 100000. | 982000. | 1800 |

# Epitome on Diametral Distance Physics 

Now therefore state the measure of the portion of the quantity of the solar forces which the planet admits into itself in any posture with respect to the sun.

We must consider the angle which the rays [radii] of the sun make with the magnetic threads of the planetary globe. For the sine of the angle complementary to this measures this portion of forces admitted. For since the efficient causes of the libration are the ray of the sun and the magnetic threads of the planet's body, two physical lines; it is right to seek the measure of the strength of the libration from the angle between these lines and from its sine.
[649] For example, let $A$ be the sun; and $I, E$, the centre of the pianetary body; $R P$ the line drawn through $A$ the sun and
 through $B$ the centre of the orbit; $E G$ and $I H$ the magnetic threads will be practically perpendicular to $R P$-at least if the counterbalancing of the semicircles is considered-and $H$ and $F$ are the solipetal termini. For it was laid down in Book iv page 95. that during the revolution of the body the magnetic threads stay practically parallel to themselves and at points $P$ and $R$ provide no occasion of attraction or repulsion, because there the positions of both the solipetal and the solifugal termini are equally distant from the sun at $A$. But in the intermediate positions, where the solipetal or the solifugal termini regard the sun directly, the strength of the libration is greatest of all. $A E$ and $A I$ are the rays of the sun. Let lines $E D$ and $I O$ be drawn parallel to line $R P$, and let perpendiculars be dropped upon them from points $F$ and $C$, wherein the rays of the sun cut the middle circles of the planetary globe; and let the perpendiculars be $C L$ and $F K$. Here the angles made by the rays of the sun with the threads are $A E G$ and $A I H$; the complementary angles are $C E D$ and $F I O$, or the arcs $C D$ and $F O$, and $C L$ and $F K$ are the sines of these angles, according as $I H$ or $E G$ is the total sine of 100,000 . Therefore it is established that as $E G$ and $I H$ are to $L C$ and $K F$, so the total quantity of the forces from the sun which are present at $I$ or $E$ is to the portion which the planet admits at the postures of the threads $E G$ and $I H$.

## Epitome on Area Rule



Then in this way let the area of the ellipse be rightly distributed between the pairs of opposite arcs; now demonstrate that the single triangles are separately the most exact measures of the single delays?

The demonstration is easy by means of the foregoing.
For since by our axiom the delay of the planet in arc $P C$ is to the delay in the equal arc $R G$, as $A P$ the distance of arc $P C$ from the source of motion is to $A R$ the distance of arc $R G$; but since also the area of triangle $P C A$ is to the area of triangle $R G A$-which has its base $R G$ equal to $P C$ the base of the first tri-angle-as $P A$ the altitude of triangle $P C A$ is to $R A$ the altitude of triangle $R G A$ : Wherefore the delay of the planet in arc $P C$ is to the delay in the equal arc $R G$ as the area of triangle $P C A$ is to the area of triangle $R G A$.

In the same way it will be demonstrated that the delay of the planet in $C F$ -which is equal in power (potestate) to $C P$-is to the delay of the same in $G H$, as the area $A C F$ is to the area $A G H$-where the sum of each pair of areas is equal to the sum of the [two] prior areas, and so on in order. Therefore the total area of the ellipse-which has been cut up at $A$ into triangles-[671] is distributed among the arcs in the same proportion wherein the total periodic time has been distributed among them. Therefore the single triangles are proportionally the most exact measures of their single arcs.
A demonstration of this full equivalence is given in my Commentaries on Mars, Chapter 59, page 291. On that page at the line Apsis longiorem, one word, erit, has brought in great obscurity; and if you change it to computaretur, everything will be clearer. Although I confess that the thing is given rather obscurely there, and most of the trouble comes from the fact that there the distances are not considered as triangles, but as numbers and lines.

# Epitome on 3/2 Power Rule Physics 

## 4. On the Causes of the Ratio of the Periodic Times

In the beginning of this consideration of movement you said that the periodic times of the planets are found to be quite exactly in the ratio of the $3 / 2$ th powers of their orbits or circles. I ask what the cause of this thing is.
Four causes come together in establishing the length of the periodic time. The first is the length of the route; the second is the weight or the amount [copia] of matter to be transported; the third is the strength of the motor virtue; the fourth is the bulk [moles] or space in which the matter to be transported is unrolled. For as is the case in a mill, where the wheel is turned by the force of the stream, so that, the wider and longer the wings, planks, or oars which you fasten to the wheel, the greater the force of the stream pouring through the width and depth which you will divert into the machine; so too
 that is the case in this celestial vortex of the solar form moving rapidly in a gyro-and this form causes the movement. [531] Consequently the more space the body- $A$ or $D$ in this case-occupies, the more widely and deeply it occupies the motor virtue, as in this case $B C A$ understood according to its width; and the more swiftly, other things being equal, is it borne forward; and the more quickly does it complete its periodic journey.

But the circular journeys of the planets are in the simple ratio of the intervals. For as $S A$ is to $S D$, so too is the whole circle $B A$ to the whole circle $E D$. But the weight, or the amount of matter in the different planets, is in the ratio of the $1 / 2 \mathrm{th}$ powers of the intervals, as was proved above, so that always the higher planet has more matter and is moved the more slowly and piles up the more time in its period, since even before now by reason of its journey it would have wanted more time. For with $S K$ taken as a mean proportional betweèn $S A$ and $S D$ the intervals of the two planets; as $S K$ is to the greater distance $S D$, so the amount of matter in the planet $A$ is to the amount in the planet $D$. But the third and fourth causes balance one another in the comparison of the different planets. But the simple ratio of the intervals plus the ratio of the $1 / 2$ th powers constitute the ratio of the $3 / 2$ th powers of the same. Therefore the periodic times are in the ratio


## Questions of "Projectability"

The evidence in Astronomia Nova shows that the four key claims - diametral distance rule, area rule, ellipse, and plane of fixed inclination through the true Sun - hold at least to high approximation for Mars over the course of its twelve revolutions of the zodiac between 1580 and 1604 .

Question: What grounds could Kepler have offered for extrapolating - i.e. "projecting" - this result beyond that period of 24 years to the motion of Mars into the indefinite past and indefinite future?

Question: What grounds could Kepler have offered for extending - i.e. projecting - this result to the motions of Jupiter, Saturn, Venus and Mercury?

In other words, to what extent and on what grounds were the chief results of Astronomia Nova directly "projectable" - to use Nelson Goodman's term into broader "nomological" generalizations?

## Eccentric Circle vs. Ellipse, Both with Area Rule


< ASQ: heliocentric longitude, circle
< ASP: heliocentric longitude, ellipse

| < E: | $\mathbf{3 0 ~ d e g}$ | $\mathbf{4 5 ~ d e g}$ | $\mathbf{6 0 ~ d e g}$ | $\mathbf{7 5 ~ d e g}$ | $\mathbf{9 0} \mathrm{deg}$ |
| :---: | :---: | :---: | :---: | :---: | :---: |
| $\mathbf{e}=\mathbf{0 . 0 2}$ |  |  |  |  |  |
| < ASQ | 29.437 | 44.201 | 59.018 | 73.899 | 88.854 |
| < ASP | 29.432 | 44.195 | 59.013 | 73.896 | 88.854 |
| diff | 17.6 sec | 20.7 sec | 18.1 sec | 11.1 sec | 0.7 sec |
| $\mathbf{e}=\mathbf{0 . 0 5}$ |  |  |  |  |  |
| < ASQ | 28.63 | 43.04 | 57.58 | 72.27 | 87.14 |
| < ASP | 28.60 | 43.01 | 57.55 | 72.25 | 87.13 |
| diff | 108.5 sec | 128.8 sec | 116.9 sec | 74.9 sec | 12.9 sec |
| $\mathbf{e}=\mathbf{0 . 0 9}$ |  |  |  |  |  |
| < ASQ | 27.61 | 41.58 | 55.73 | 70.14 | 84.86 |
| < ASP | 27.51 | 41.46 | 55.63 | 70.01 | 84.84 |
| diff | 5.73 min | 6.94 min | 6.51 min | 4.47 min | 1.25 min |

For purposes of comparison, the eccentricity of the Earth is a little less than 0.02 ; the eccentricity of Jupiter a little less and that of Saturn a little greater than 0.05 ; and the eccentricity of Mars a little greater than 0.09 . In short, without Tycho's data, the distinction between circle and ellipse would not have been detectable, and even with it this difference was not detectable for Venus, Earth, Jupiter and Saturn.

Finally, a circle with equant and bisected eccentricity gives, for $\mathrm{e}=0.09$ and $\mathrm{E}=45 \mathrm{deg}$, an angle ASP' of 41.36 deg and hence a difference between the angle ASP of the ellipse plus area rule of 5.84 min . So, asking astronomers before Tycho to have noticed the difference between a circle with bisected eccentricity and an equant and a Keplerian ellipse is asking for something they could not have done.

# RUDOLPHIN Æ， 

உUIBUSASTRONOMICASCIENTIA，TEMPO－ rum longinquitate collapfa Restauratio continetur；

A Phœnice illo Aftronomorum
T Y C H O N E

Ex Illuffri \＆r Generofa Brat eor rum in Regno Danic
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PRIMUM ANIMOCONCEPTA ET DESTINATA ANNO CHRISTI MDLXIV：EXINDE OBSERVATIONIBUS SIDERUM ACCURA－ TISSIMIS，POST ANNUM PRECIPUE MDLXXII，QUO SIDUS IN CASSIOPEJ届 Constellatione notumeffulsit，serid affectata；varitsque operibus，cum mb－ chanicis，tùmlibrariis，impenfo parrimoniv annplifimo，accedentibus etiam fubfidiis Friderici il．Danie REGIS，regali magnificentiadignis，tracta per annos XXV，potifimùm in Infula frerı Sundici HuEN－ na，\＆arce Uraniburgo，in hosufusà fundamentis extructà：

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tabulas ipsas，Jam et nuncupatas，et affectas，sed morte Authoris sut anno MDCI desertas，

## FUSSU ETSTIPENDIIS FRETUS TRIUM IMPPP．

## RUDOLPHI，MATTHI E，FERDINANDI，

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Anna M．D C．XXVIL．


The Keplerian Ellipse


$$
\begin{aligned}
C A & =\text { mean durance } \\
& =a
\end{aligned}
$$

$\frac{C S}{C A}=$ eccentricity

$$
=e
$$

start with the eccentricity, the length of line of apsides (ZCA), and its orientation, as given by the (heliocentric) longitude of its aphelion


Next draw a circle of radws $r=C A$, with the linic of apsides as a diameter. Then, for any angle $A C Q=E$, the so-called "eccentric anomaly," the location of the planet on the elliptic trajectory is given as follows


$$
\begin{aligned}
S P & =a(1+e \cos E) \\
\angle A S P & =\sin ^{-1}\left(\frac{a \sqrt{1-e^{2}} \sin E}{S P}\right) \\
& =\sin ^{-1}\left(\frac{\sqrt{1-e^{2}} \sin E}{1+e \cos E}\right)
\end{aligned}
$$

The elliptic trajectory is then given by the
locus of points $P$ (as E ́goes from 0 to 360 dey)

## Elements of Keplerian Orbits

Orientation of the line of apsides - i.e. longitude of aphelion: $\omega$

Length of the semi-major axis - i.e. mean distance of planet from Sun: $a$

Eccentricity: $e$
Period: $\boldsymbol{P}$ - or mean motion in longitude $\boldsymbol{n}$
Orientation of the line of nodes - i.e. longitude of ascending node: $\Omega$

Inclination: $\boldsymbol{i}$

Mean longitude at epoch: $L$, or time last at aphelion, $T$

## An Assessment of Kepler's Orbital Elements

Based on Those Inferred for 1600 from Newcomb

|  |  | Eccentricity | Apehelion | Mean Distance |
| :--- | :--- | :--- | :--- | :--- |
| Mercury | N | 0.20555 | $251^{\circ} 14^{\prime} 09^{\prime \prime}$ | 0.38710 |
|  | K | 0.21005 | $252^{\circ} 49^{\prime} 58^{\prime \prime}$ | 0.38808 |
| Venus | N | 0.00697 | $305^{\circ} 55^{\prime} 51^{\prime \prime}$ |  |
|  | K | 0.00692 | $301^{\circ} 14^{\prime} 22^{\prime \prime}$ | 0.72333 |
|  |  | N | 0.01688 | $276^{\circ} 04^{\prime} 02^{\prime \prime}$ |




## Numerical Values


 different values appear in other references (and even within the same reference). In some instances there are "official" values, so that different computations share a common set, making comparison possible. The following values are included partly for their interest, and partly so that you can include them, if you wish, in your computations.
C. 1 Orbital Elements of Planets

The data for the planets (excluding Pluto) were kindly supplied by L. E. Doggett from data of P. Bretagnon (Reference 72). They are mean elements, referred to the equinox J 2000.0 . $T$ is measured in Julian centuries from that epoch. The semimajor axes have been derived from the terms in $L$ (the mean longitude at epoch) factored by $T$, with allowance first made for the precession in longitude. (See Appendix C7.)

For detailed information on osculating elements, refer to the Astronomical Almanac. For low-precision formulas, see Reference 38. ૪ $\tilde{\omega}=77.456119+1.556477 \times T+0.000296 \times T^{2}$ $\Omega=48.330893+1.186188 \times T+0.000176 \times T^{2}$ $i=7.004986+0.001821 \times T-0.000018 \times T^{2}$ $e=0.20563175+0.00002041 \times T-0.00000003 \times T^{2}$ $L=252.250906+149474.072249 \times T+0.000304 \times T^{2}$ $a=0.38710353$

## Appendix C



$$
\begin{aligned}
C A & =a \\
\frac{C S}{C A} & =e
\end{aligned}
$$

$$
\text { fraction of period }=\frac{n(t-T)}{2 \pi}=\frac{M}{2 \pi}
$$

where $n$ is mean (angular) motion, $T$ is time of aphelion, and $M$ is called "mean anomaly"

$$
\begin{aligned}
\frac{M}{2 \pi}=\frac{\text { area of sector } A S P}{\text { area of ellipse }} & =\frac{\text { area of sector } A S Q}{\text { area of circle }} \\
& =\frac{\frac{a^{2}}{2} E+\frac{a^{2} e}{2} \sin E}{\pi a^{2}} \\
& =\frac{E+e \sin E}{2 \pi}
\end{aligned}
$$

Therefore

$$
n(t-T)=M=E+e \sin E \quad \text { (Kepler's equation) }
$$

Where $E$ is called the "eccentric anomaly"


You have shown how to compute the mean anomaly and the coequated (true) anomaly from a given eccentric anomaly. But more frequently use requires that for a given mean anomaly, as from a given time, we find the others. Explain that also.

Here there is no direct way, but one who wishes to compute this without tables must employ the 'rule of assumptions', namely, as shown in the following figure, assuming the eccentric anomaly (arc) $A P^{\prime}$ as such and such an amount, and for the eccentric anomaly so assumed, computing its mean anomaly (area) $A S P^{\prime}$. And if the result is the amount (of the mean anomaly) that was proposed, the eccentric anomaly (arc) $A P^{\prime}$ will have been assumed correctly. But if the result is not such an amount, the assumption will have to be corrected by means of the result and the work repeated.

## Kepler's "Laws"

1. (The) planets describe elliptical trajectories about the Sun (or central body) at a focus
2. The radius vectors of (the) planets sweep out equal areas in equal times about the Sun (or central body)
3. The periods of (the) planets are in sesquialternate proportion to their mean distances from the Sun (or central body)
4. The distances of (the) planets from the Sun (or central body) vary as $a(1+e \cos E)$, where $a$ is the mean distance, $e$ is the eccentricity, and $E$ is the "eccentric anomaly"
5. The trajectories defined by (the) planets are confined to single planes through the Sun (or central body) at a fixed inclination to the plane of the ecliptic

## Keplerian Motion: Projectability

At least to high approximation, the five planets move along ellipses, sweeping out equal areas in equal times with respect to the (true) Sun located at a focus common to all, on planes passing through the Sun at fixed angles of inclination, in periods proportional to the $3 / 2$ power of their mean distances from the Sun.

## Questions about this group of generalizations:

1. What grounds were there for extending - i.e. projecting - each of the generalizations beyond the five planets to support claims about:
a. Any ("possible") body orbiting the Sun?
b. Any body engaged in celestial orbital motion - e.g. Moon, satellites of planets?
c. Any celestial body moving within our planetary system - e.g. comets?
2. What grounds were there for concluding that the specific statement of each generalization was properly suited for any such projection?
3. What, if any, further qualifications - e.g. tacit ceteris paribus conditions - needed to be noted with each generalization before projecting it?

## Status of the Generalizations

1) Exact, essentially exact, merely approximate
2) If not exact, then which combination of:
a. Holds in the mean or not in the mean (skewed) - e.g. ideal gas law does not hold in the mean
b. Idealization or mere approximation
where idealizations can be:
i. Generalizations that would hold exactly in certain specifiable circumstances
ii. Generalizations that serve only to provide simpler mathematics - e.g. linear elasticity

Question: Assuming, as Kepler had by 1625, that his generalizations are probably not exact, what, if any evidence could have indicated for each which of the above categories it belonged in?

# Kepler's Three Lines of Evidential Reasoning 

## Astronomia Nova (1609)

Start from Tycho's observations and use past practices in mathematical astronomy, including most notably Tycho's solar theory, to leverage reforms in those practices and a new orbit for Mars, invoking plausible physical considerations and precision of inferences from observation to conclude that the reforms and new orbit amount to something more than merely a representation of Tycho's data to (surprisingly) high approximation; thereby justify the claim that, absent the emergence of any substantial discrepancies with observation, the reforms and hence the orbital schema for Mars hold for the other planets as well.

## Epitome Astronomiae Copernicanae (1618-21)

Start from proposed physical principles, defended in part through analogy with known physics and in part by anticipatory claims about their consequences agreeing with observation, to derive, with minimal auxiliary assumptions, the reforms and orbital schema as holding exactly (where no exteernal, secondary physical effects involved) for all the planets; and then (attempt to) show that the far more complicated motion of the Moon, theretofore never successfully represented to within observational precision, can be derived from the same physical principles without need of further auxiliary assumptions.

## Tabulae Rudolphinae (1627)

Start from the reforms and orbital schema, taken as given, and use appropriate observations to derive values of the elements for each planet, opening the way to comparisons between calculation and all available observations, with provision for refining the values of the elements on the basis of these and further observations; success in reducing the discrepancies between calculation and observation through such refinements is evidence that the reforms and orbital schema amount to something more than merely one among many possible representations of Tycho's data to high approximation.

## Fundamental Evidence Problems in Orbital Astronomy

1. How to determine the distances of celestial objects from the Earth -- in particular, the distance of Mercury, Venus, Mars, Jupiter, and Saturn as well as the Sun and Moon (and even comets) -- not only on average, but also over the course of time.
2. How to determine which apparent celestial motions and changes of motion, as observed from the Earth, are merely apparent and which are real, at least relative to the fixed stars.
3. How to determine whether such calculational rules as Kepler's area rule and ellipse, strikingly accurate as they seem to be, are merely one among many comparably accurate, though still only approximate, ways of representing Tycho's observations, or instead are something more than this.
i. Whether they hold exactly or at least would hold exactly were it not for various "external" factors.
ii. Whether they project beyond the given observations indefinitely into the past for the celestial objects in question and whether they project beyond these objects to other possible celestial objects.
4. How to determine what physical principles are governing the motions of the various celestial objects -- especially the planets, but also the Sun and Moon and maybe comets as well.

The first two of these problems go back at least as far as Ptolemy. The third also could be asked of Ptolemy's calculational rules, though asking it becomes much more appropriate after the factor of 50 or so improvement in accuracy achieved by Kepler. The fourth was asked in ancient times as well, though answered largely by philosophical appeals to the perfection of circular motion. Taken as requesting empirical evidence, it comes to the fore more strongly after Kepler turns to physical principles as the appropriate way of settling the other three evidence problems.

## SIX ASPECTS OF THE EVIDENCE PROBLEM IN

## PLANETARY ASTRONOMY RECOGNIZED BY KEPLER

Whether Kepler should be regarded as the first to take the problem of planetary astronomy to be that of establishing a physically correct account of orbital motion is unclear. Nevertheless, he was almost certainly the first to realize the magnitude of this problem, for he did appreciate the challenge posed by the following issues:

1. The possibility not just of systematic error in the raw data, but of the corrections introduced to compensate for systematic error being inadequate.
2. The possibility of being led down a garden path by the need to use theoretical assumptions of some sort or other to extract conclusions from astronomical data, returning only much later to assess and, if need be, to refine these assumptions.
3. The fact that any number of distinct trajectories can, in principle, agree with imprecise data to more or less the same level of accuracy -- no matter how small the imprecision is -- so that the choice of any one of these trajectories over the others cannot be dictated by astronomical data alone.
4. The possibility of being led down a garden path by the need to appeal to underlying physics in one way or another in reaching conclusions about the physically true trajectory, while at the same time having to use conclusions about this trajectory as the principal basis for evidence for the conjectures about the underlying physics.
5. How to decide whether residual discrepancies between theory and observation should be dismissed, at least for the moment, as arising from observational error or extraneous secondary physical effects, or whether they should instead be regarded as falsifying the theory.
6. How to reach conclusions -- especially in the face of the above difficulties, but even without them -- about orbital motion in the remote past and in the long term future and about whether other orbiting bodies would have the same trajectories as those now accessible to investigation.

## Sources for the Material in the Handout

For the chronology and attached figures, James Voelkel, Johannes Kepler and the New Astronomy, Oxford University Press, 1999.

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For the comparison of Kepler's values of the elements and those for the same time inferred from Newcomb, Curtis Wilson, "Predictive astronomy in the century after Kepler," in R. Taton and C. Wilson, Planetary Astronomy from the Renaissance to the Rise of Astrophysics, Part A: Tycho Brahe to Newton, Cambridge University Press, 1989, p. 180.

